



A New Broadening Technique of the Numerically Unresolved Solar Transition Region and Its Effect on the Spectroscopic Synthesis Using Coronal Approximation

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Abstract

The transition region is a thin layer of the solar atmosphere that controls the energy loss from the solar corona. Large numbers of grid points are required to resolve this thin transition region fully in numerical modeling. In this study, we propose a new numerical treatment, called LTRAC, which can be easily extended to the multidimensional domains. We have tested the proposed method using a one-dimensional hydrodynamic model of a coronal loop in an active region. The LTRAC method enables modeling of the transition region with a numerical grid size of 50–100 km, which is about 1000 times larger than the physically required value. We used the velocity differential emission measure to evaluate the possible effects on the optically thin emission. Lower-temperature emissions were better reproduced by the LTRAC method than by previous methods. Doppler shift and nonthermal width of the synthesized line emission agree with those from a high-resolution reference simulation within an error of several kilometers per second above the formation temperature of 10^5 K.

Unified Astronomy Thesaurus concepts: [Solar transition region \(1532\)](#); [Solar corona \(1483\)](#); [Hydrodynamical simulations \(767\)](#)

1. Introduction

The solar transition region is a thin layer between the chromosphere and the corona. In this region, the temperature and mass density change steeply from chromospheric values to coronal values. The rapid change of the mass density produces the steep gradients of the acoustic and Alfvén speeds, which leads to wave reflection. The rate of reflection of Alfvén waves at the thin transition region is roughly determined by the density variation in that region (Hollweg 1984; Verdini et al. 2012). The free magnetic energy injected into the solar corona is released through impulsive heating events in both the magnetic braiding and the wave heating models (e.g., van Ballegoijen et al. 2011; Rempel 2017; Matsumoto 2018). The thermal energy in the solar corona is transported by thermal conduction down to the transition region, releasing the energy as radiation into space. The resulting evaporation of the chromospheric plasma increases the coronal density (Reale 2010; Klimchuk 2015). Thus, accurate modeling of the transition region cannot be avoided in modeling the energy exchange between the chromosphere and the corona.

A simple static, coronal loop model was introduced by Rosner et al. (1978). The major assumptions in this model are as follows: (1) a coronal loop is sufficiently short so that the spatial variation of the gas pressure can be ignored, (2) a volumetric heating rate in a coronal loop is spatially uniform, and (3) the energy balance in a static loop is given by

$$\frac{\partial}{\partial s} \left(\kappa \frac{\partial T}{\partial s} \right) + Q_{\text{ext}} + Q_{\text{rad}} = 0, \quad (1)$$

where T is the temperature, s is the coordinate along the coronal loop, κ is the thermal conduction coefficient, Q_{ext} is the volumetric heating rate, and Q_{rad} is the optically thin radiative cooling rate. In this model, the temperature profile near the transition region is sensitive to the gas pressure (or heating rate) but is not sensitive to the length of the loop. We estimated the thickness of the transition region using the spatial profile of the

temperature given by Equation (C1) in Appendix C of Rosner et al. (1978). The thickness of the transition region was defined as a distance $s(T_1) - s(T_0)$, where $T_0 = 2 \times 10^4$ K (i.e., the temperature at the footpoint of the coronal loop), $T_1 = 10^5$ K (the typical temperature at the transition region), and $s(T)$ is the coordinate s at the temperature T . The estimated thickness of the transition region was approximately $s(T_1) - s(T_0) \sim 28 \times (T_{\text{top}}[\text{K}] / 10^6)^{-3}$ km, where T_{top} is the temperature at the apex of the coronal loop. Even the quiet corona cannot be resolved by a grid size of several tens of kilometers, which is typically assumed in multidimensional simulations (e.g., Abnett 2007; Gudiksen et al. 2011; Iijima & Yokoyama 2015, 2017; Rempel 2017). The transition region becomes even thinner in hotter active regions or flaring loops owing to the strong temperature sensitivity of the thermal conduction coefficient and the radiative cooling.

The lack of sufficient resolution of the transition region in hydrodynamic and magnetohydrodynamic simulations has been a topic of continuing discussion (Craig et al. 1982; Klimchuk 2006; Bradshaw & Cargill 2013). One possible solution is to use a moving mesh (Carlsson & Stein 1992) or adaptive mesh refinement (Bradshaw & Klimchuk 2011). However, it is difficult to achieve such small numerical grid sizes in multidimensional simulations owing to the resulting high computational cost. As the coronal heating processes (both magnetic braiding and wave heating) are intrinsically three-dimensional in nature, it is necessary to devise a way to capture the transition region adequately on a coarse numerical grid to model the solar corona realistically.

A technique for broadening the numerical transition region has been suggested by Linker et al. (2001), Lionello et al. (2009), and Mikić et al. (2013). By introducing the problem-dependent parameter T_c , they suggested that the thermal conduction coefficient should be enhanced by a factor $f_c = \max(1, T/T_c)$ to broaden the unresolved transition region. They also reduced the radiative cooling rate by the same factor

f_c so that the total radiative loss remains unchanged. Johnston & Bradshaw (2019) and Johnston et al. (2020) suggested a way to model the transition using a method they term “transition region adaptive conduction” (TRAC). This method automatically determines the optimal value of the parameter T_c in one-dimensional coronal loop simulations. By selecting the highest temperature at a location with a high-temperature gradient, they showed that the dynamic evolution of a coronal loop can be reproduced with the coarse grid size of 100 km. However, a multidimensional extension of the TRAC method is not straightforward, because we need a way to track the magnetic field line as the optimal value of T_c is determined along the field line (Zhou et al. 2021). Johnston et al. (2017a, 2017b) also suggested an alternative way to resolve the transition region using jump conditions.

Spectroscopic observations in the EUV and X-ray are important diagnostic tools for the solar corona and transition region, especially for information about plasma motions (e.g., Chae et al. 1998; Teriaca et al. 1999; Doschek et al. 2007; Imada et al. 2008). In particular, the statistical characteristics of nonthermal line widths and the Doppler velocities have been used as constraints on coronal heating models (e.g., Peter & Judge 1999; Imada et al. 2009; Brooks & Warren 2016; Testa et al. 2016; van Ballegoijen et al. 2017). As the line profiles contain considerable information about the coronal plasma, a numerical treatment of the broadening method must therefore reproduce the spectroscopic observables.

In this study, we propose a new approach called the LTRAC (localized TRAC) method to broaden the unresolved transition region in numerical simulations. The local nature of the LTRAC method allows us to apply it to the parallel computation of multidimensional simulations using the domain decomposition technique. As a large number of grid points is required to resolve the transition region properly, we tested the proposed method using a one-dimensional hydrodynamic model of a solar coronal loop. We investigated the performance of the proposed method by examining the time evolution of the plasma in the transition region and corona, specifically focusing on spectroscopic observables in the optically thin approximation. We found that the LTRAC method reproduces the spectroscopic observables coming from the low-temperature plasma better than previous methods.

2. Method

2.1. Simulation Setup

To investigate the performance of the methods used to broaden the transition region, we solved the one-dimensional hydrodynamic equations along a field line, including the effects of gravity, Spitzer–Härm conduction, radiative cooling, and external heating. The basic equations are the continuity equation

$$\frac{\partial \rho}{\partial t} + \frac{\partial}{\partial s}(\rho V_s) = 0, \quad (2)$$

the equation of motion

$$\frac{\partial}{\partial t}(\rho V_s) + \frac{\partial}{\partial s}(\rho V_s^2) + \frac{\partial P}{\partial s} = \rho g_s, \quad (3)$$

and the internal energy equation

$$\begin{aligned} \frac{\partial e}{\partial t} + \frac{\partial}{\partial s}[(e + P)V_s] \\ = V_s \frac{\partial P}{\partial s} + Q_{\text{cnd}} + Q_{\text{rad}} + Q_{\text{ext}}, \end{aligned} \quad (4)$$

where ρ is the mass density, P is the gas pressure, e is the internal energy density, V_s is the velocity along a field line, and g_s is the gravitational acceleration along a field line. The equation of state is computed assuming local thermodynamic equilibrium and considering the most abundant six elements in the solar atmosphere. The ionization energies of the six elements and the latent internal energy of molecular hydrogen are included in the internal energy of the plasma.

The quantity Q_{cnd} represents the volumetric heating rate due to the thermal conduction:

$$Q_{\text{cnd}} = \frac{\partial}{\partial s} \left(\kappa \frac{\partial T}{\partial s} \right). \quad (5)$$

For the thermal conduction coefficient, we assume Spitzer–Härm conduction $\kappa = \kappa_0 T^{5/2}$ with $\kappa_0 = 10^{-6}$ in cgs units.

The radiative cooling rate Q_{rad} is a combination of the cooling rates from the optically thick and thin regions, as given by

$$Q_{\text{rad}} = (1 - \xi)Q_{\text{thick}} + \xi Q_{\text{thin}}, \quad (6)$$

where $\xi = \exp(-P/P_{\text{thick}})$ is the switching function between optically thick cooling, Q_{thick} , and thin cooling, Q_{thin} . We chose $P_{\text{thick}} = 10^4$ dyn cm $^{-2}$ as the threshold parameter. The optically thick cooling rate is given by

$$Q_{\text{thick}} = -\rho C_V \frac{T - T_{\text{rad}}}{\tau}, \quad (7)$$

where C_V is the isochoric heat capacity per unit mass and $T_{\text{rad}} = 6000$ K is the radiation temperature. Following the approach of Gudiksen & Nordlund (2005), we chose the cooling timescale τ to be

$$\tau = \max \left(0.1 \sqrt{\frac{\rho_{\text{surf}}}{\rho}}, 4\Delta t \right), \quad (8)$$

where $\rho_{\text{surf}} = 2 \times 10^{-7}$ g cm $^{-3}$. We used the slightly modified version of the optically thin cooling rate given by

$$Q_{\text{thin}} = -n_e n_H \Lambda(T) \left[1 - \left(\frac{T_{\text{rad}}}{T} \right)^4 \right] \quad (9)$$

to prevent the formation of low-temperature regions with $T < T_{\text{rad}}$. Here, n_e and n_H represent the number densities of electrons and hydrogen nuclei. This modification does not affect either the transition region or the corona.

The numerical scheme is based on the finite-difference method. The hyperbolic equations are solved using the second-order SLIP scheme (Jameson 1995) with the third-order strongly stability preserving Runge–Kutta method (Shu & Osher 1988). The energy-consistent formulation by Iijima (2021) is used in the spatial discretization to ensure the total energy conservation while explicitly solving the internal energy equation. We used the operator splitting between the hydrodynamic equations and the thermal conduction equation. The thermal conduction equation is discretized using the

second-order central difference method, and it is integrated implicitly using the backward Euler method to avoid the severe time-step restriction encountered in explicit schemes.

The numerical domain extends from the solar surface ($s = 0$) to the top of the coronal loop ($s = L$). The half length of the coronal loop L is chosen to be 25 Mm in a typical case. We assumed a semicircular coronal loop and reduced the field-aligned gravitational acceleration g , following the geometrical effect. We calculated the initial temperature profile from

$$\ln T(s) = \ln T_{\text{surf}} + \frac{1}{2} \left[1 + \tanh \left(\frac{s - s_{\text{TR}}}{\Delta s_{\text{TR}}} \right) \right] (\ln T_{\text{surf}} - \ln T_{\text{top}}), \quad (10)$$

where $s_{\text{TR}} = 2.5$ Mm, $\Delta s_{\text{TR}} = 500$ km, $T_{\text{surf}} = 6000$ K, and $T_{\text{top}} = 10^6$ K. The mass density is integrated assuming hydrostatic equilibrium, with the boundary value of $\rho = 2 \times 10^{-7}$ g cm $^{-3}$. This initial condition is not energetically balanced because of the thermal conduction, radiative cooling, and external heating terms. We assume reflective boundary conditions at the top and bottom of the domain. In this study, the dynamical evolution of the simulated corona is controlled by imposing an external heating rate Q_{ext} to mimic coronal heating due to arbitrary processes. In a typical case, we impose spatially uniform volumetric heating rate $Q_{\text{ext}} = F_{\text{ext}}/L$. The injected energy flux F_{ext} is set to 3×10^7 erg cm $^{-2}$ s $^{-1}$ for the initial 600 s of the simulation and set to zero after the time of 600 s. This setup mimics a coronal loop in an active region.

2.2. The LTRAC Method

We now describe the LTRAC method we implemented to broaden the unresolved transition region in the model described in Section 2.1. As our new method is an extension of the methods suggested by Lionello et al. (2009) and Johnston & Bradshaw (2019), we summarize their methods first.

Lionello et al. (2009) used energy balance between the radiative cooling and conductive heating,

$$Q_{\text{rad}} \sim \frac{\kappa T}{\delta_{\text{TR}}^2}, \quad (11)$$

to obtain the approximate thickness of the transition region as

$$\delta_{\text{TR}} \sim \sqrt{\frac{\kappa T}{Q_{\text{rad}}}}. \quad (12)$$

Here, we note that if we assume that the typical temperature in the transition region does not depend on the temperature at the loop top, the temperature dependence of the transition region thickness (discussed in Section 1) can be derived as $\delta_{\text{TR}} \propto T_{\text{top}}^{-3/2}$, where we used the scaling law of Equation (4.3) in Rosner et al. (1978). The dependence on the loop top temperature T_{top} comes from the dependence on the mass density of the radiative cooling rate Q_{rad} .

Lionello et al. (2009) proposed that if κ and Q_{rad} are substituted by $f_c \kappa$ and Q_{rad}/f_c , respectively, the transition region can be broadened by a factor of f_c . This substitution can be alternatively expressed by

$$\kappa^{\text{b}} = f_c \kappa \text{ and } Q_{\text{rad}}^{\text{b}} = Q_{\text{rad}}/f_c, \quad (13)$$

where the variable with superscript U^{b} indicates the variable U modified by the broadening technique U of the transition region. The substitution (13) broadens the approximate thickness of the transition region (12) as $\delta_{\text{TR}}^{\text{b}} = f_c \delta_{\text{TR}}$. We did not include the effect of the heat flux saturation (Fisher et al. 1985; Patsourakos & Klimchuk 2005; Bradshaw & Cargill 2006) for simplicity, which is important for very hot coronal loops. One possible way to implement the heat flux saturation is to enhance the saturated (limited) heat flux (or conductive flux) by a factor of f_c instead of the thermal conduction coefficient κ . However, its validity and performance should be investigated in the future study.

Johnston et al. (2020) pointed out that the substitution (13) does not alter the total radiative loss from the broadened transition region. From Equation (12), the total radiative loss integrated across the transition region can be approximated as

$$\int_{s_0}^{s_1} Q_{\text{rad}} ds \sim Q_{\text{rad}} \delta_{\text{TR}} \sim \sqrt{\kappa Q_{\text{rad}} T} = \sqrt{\kappa^{\text{b}} Q_{\text{rad}}^{\text{b}} T}, \quad (14)$$

where s_0 and s_1 are the spatial coordinates at the bottom and top of the transition region, respectively, and we used Equation (13) to derive the last equality. As the substitution of κ and Q_{rad} by Equation (13) does not alter the right-hand side of Equation (14), the total radiative loss from the transition region is not altered even if the transition region is numerically broadened. Similarly, the total conductive heating of the transition region ($\kappa T/\delta_{\text{TR}}$) remains unchanged by the substitution (13).

Lionello et al. (2009) suggested to determine the broadening factor f_c as

$$f_c = \max(1, T_c/T)^{5/2}, \quad (15)$$

where T_c is a parameter that depends on the specific problem and the grid size. For example, T_c is set to a typical temperature at the top of the transition region. Lionello et al. (2009) noted that the lower T_c requires higher numerical resolution to resolve the transition region. Johnston & Bradshaw (2019) proposed the TRAC method that determines the optimal value of T_c from

$$T_c = \max_s (T(s)) \text{ s.t. } \left| \frac{\Delta s}{T} \frac{\partial T}{\partial s} \right| > \sigma_c, \quad (16)$$

where Δs is the grid size and the problem-independent parameter $\sigma_c = 1/2$ is introduced. Following Johnston & Bradshaw (2019), the variable T_c is limited by imposing a lower bound of typical chromospheric temperature (2×10^4 K) and an upper limit of 20% of the maximum coronal temperature. Note that a multi-dimensional extension of the TRAC method is not straightforward, as the maximum is taken along a field line. In parallel simulations with the domain decomposition, the computation of the maximum value along a field line requires global communication, which may result in less parallel efficiency.

There are slight differences between the implementations of the TRAC method in Johnston & Bradshaw (2019) and Johnston et al. (2020). Our implementation of the TRAC method is based on Johnston & Bradshaw (2019), not Johnston et al. (2020). Johnston et al. (2020) suggested to divide the external heating rate by the factor f_c . We do not follow their approach because the modification of the external heating rate may cause an error in the total energy injected into the whole atmosphere. As we assume a spatially uniform volumetric

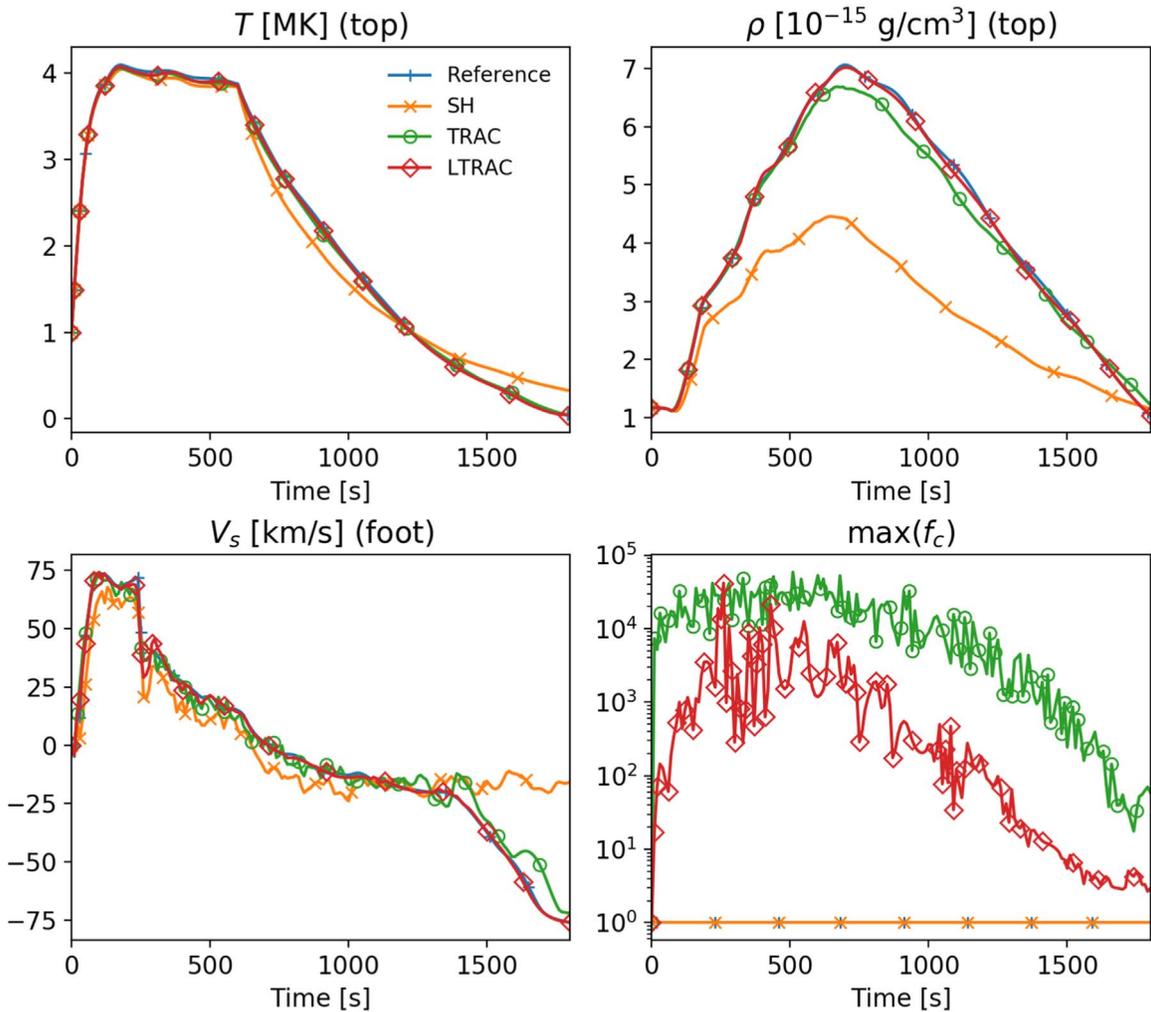


Figure 1. Time variation of the temperature at the loop top (top left), the mass density at the loop top (top right), the velocity at the footpoint (bottom left), and the maximum value of the transition region broadening factor f_c (bottom right). The coronal loop is heated using a spatially uniform volumetric heating term from $t = 0$ s to $t = 600$ s (see Section 2.1 for details). Shown are the numerical solutions calculated using the SH method (orange line with crosses), using the TRAC method (green line with circles), and using the LTRAC method (red line with diamonds). The reference solution is shown as a blue line with plus signs. A uniform grid size of 50 km is used except for the reference.

heating rate, this minor difference in the heating term does not affect the results presented in this paper. Johnston et al. (2020) also suggested to use the temporal smoothing on T_c . This modification may reduce the temporal oscillation of the low-temperature emission measure in the TRAC method observed in Figure 4. In this study, we did not use this temporal smoothing to clarify the effect of the different spatial profiles of the broadening factor between the TRAC and LTRAC methods.

In this study, we propose a new method called LTRAC to determine the broadening factor f_c using only the information from nearby grid points. Similar to the TRAC method, the LTRAC method uses the temperature gradient to detect the unresolved transition region. We define the normalized temperature gradient as

$$r = \exp \left| \frac{\Delta s}{T} \frac{\partial T}{\partial s} \right|. \quad (17)$$

In the unresolved transition region, the quantity r becomes large and the nonnegligible numerical artifacts can lead to unphysical solutions (e.g., the wrong density in the corona; see also Figures 1

and 2). To reduce the steep temperature gradient in the high r -value region, the broadening factor f_c is determined as

$$f_c = \overline{\max(1, r/r_c)^p}, \quad (18)$$

where r_c and p are nondimensional parameters and the overline indicates a low-pass filter in space. The threshold parameter $r_c > 1$ is introduced so that the broadening factor f_c is kept at unity (no broadening) in the grid points with a smooth temperature profile of $r < r_c$. The parameter $p \geq 1$ is introduced to accelerate convergence. We used $r_c = 1.5$ and $p = 4$ as typical values in this study. The dependence of the LTRAC method on these parameters is investigated in Appendix B. The low-pass filter is introduced to smooth the spatial distribution of f_c computed from the numerical gradient of the temperature, which is generally nonsmooth. As our discretization described in Section 2.1 is second order in space, we used the second-order low-pass filter defined by

$$\overline{U}_j = \frac{U_{j-1} + 2U_j + U_{j+1}}{4}, \quad (19)$$

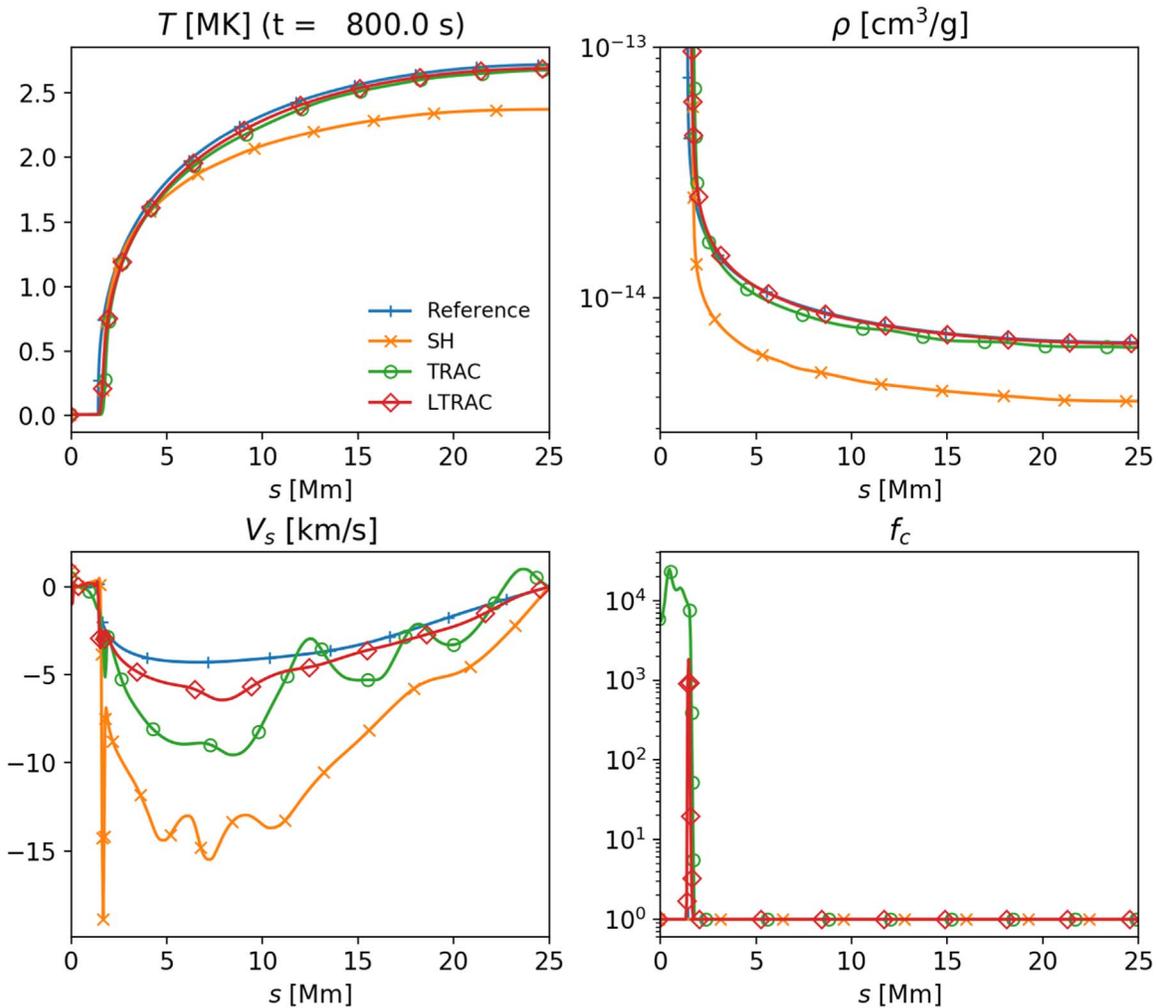


Figure 2. Spatial variation along a coronal loop at $t = 800$ s. Shown are the temperature (top left), the mass density (top right), the velocity (bottom left), and the transition region broadening factor f_c (bottom right). The notation is the same as that in Figure 1.

where U_j indicates the value of a variable U at the j th grid point along the s -coordinate. We note that the procedure above can be extended easily into multidimensional geometries, which is one of the advantages of the LTRAC method.

Here, we summarize the broadening process of the transition region by the LTRAC method. When the temperature gradient becomes steeper ($r > r_c$), f_c is enhanced by Equation (18). The enhanced f_c broadens the transition region through the enhanced thermal conduction by Equation (13) while preserving the total radiative loss (Equation (14)). In the broadened transition region, the broadening factor f_c decreases and approaches unity. The value of f_c in the steady state is determined by the balance between the broadening effect by the LTRAC method and the thinning effect by the physical nature of the transition region.

2.3. Velocity Differential Emission Measure

To investigate the effect on the spectroscopic observables, we use the velocity differential emission measure (VDEM) for optically thin emissions. The VDEM is defined by

$$\int_T \left[\int_V \text{VDEM}(T, V) dV \right] dT = \int_h n_e n_H dh, \quad (20)$$

where V is the line-of-sight velocity and the integration on the right-hand side is taken along the line of sight. Let us focus on the optically thin intensity that can be approximated by

$$I(\lambda) = \int_h C(T, V, \lambda) n_e n_H dh. \quad (21)$$

Equation (21) is usually valid if ionization/excitation equilibrium is satisfied in the coronal approximation (Del Zanna & Mason 2018). The VDEM can fully reproduce the intensity, as shown by

$$I(\lambda) = \int_T \left[\int_V C(T, V, \lambda) \text{VDEM}(T, V) dV \right] dT. \quad (22)$$

The effect of the proposed method on the optically thin spectroscopic observables can therefore be analyzed using the VDEM. The zeroth-, first-, and second-order moments of VDEM with respect to V correspond to the usual differential emission measure (DEM), Doppler velocity, and nonthermal line width, respectively. The definition of the VDEM given above is identical to that given by Cheung et al. (2019), but it is different from the original version presented by Newton et al. (1995).

We found that calculations of the optically thin emission that employ the methods for broadening the unresolved transition region require special care. As both the TRAC and LTRAC methods use the substitution (13), the local radiative cooling rate is reduced by a factor of f_c so that the total radiative loss remains unchanged by the numerical broadening of the transition region (see Equation (14)). Similarly, the optically thin emission from the numerical solution must be calculated taking the broadening factor into account:

$$I(\lambda) = \int_h C(T, V, \lambda) n_e n_H \frac{dh}{f_c}. \quad (23)$$

The VDEM must be computed in a similar manner. This correction mainly affects the low-temperature region.

3. Results

In this section, we analyze the performance of the LTRAC method compared with previous methods. We investigated three different methods for treating the unresolved transition region: one uses the classical Spitzer–Härm conduction without applying any broadening technique (denoted as SH for brevity), one uses the TRAC method, and one uses the LTRAC method.

3.1. Overall Structure

The time variation of the spatially averaged variables calculated with a uniform grid size of 50 km is shown in Figure 1. The variables at the loop top are averaged over the interval of $0.6L \leq s \leq L$, whereas the variables at the footpoint are averaged over $4 \text{ Mm} \leq s \leq 6 \text{ Mm}$. The high-resolution reference solution is calculated using the SH method with a grid size near the transition region of 100 m. We refer the reader to Appendix A for the details of the reference solution. The temperature of the underresolved SH solution (with the grid size of 50 km) slightly deviates from the reference solution. A larger deviation of the underresolved SH solution from the reference solution can be observed in the mass density (Bradshaw & Cargill 2013). The mass density at the loop top is more than 30% smaller than that of the reference solution. In contrast, the numerical solutions obtained using the TRAC or LTRAC methods show good agreement with the reference solution. For the variation of the mass density and velocity, the LTRAC solution is slightly close to the reference solution. The maximum value of the broadening factor f_c in the LTRAC solution is about one order of magnitude smaller than that in the TRAC solution. The smaller value of f_c implies smaller thermal conduction coefficients, which is advantageous if the thermal conduction term is integrated explicitly.

The maximum value of the broadening factor f_c in the LTRAC method shows temporal oscillation as shown in Figure 1. We could not fully understand its origin, but the oscillatory nature may be caused by a kind of feedback between the broadening effect of the transition region by the LTRAC method and the thinning effect from the nature of the transition region. For a better understanding, we checked the dependence on the size of time stepping by running the simulation with 10 times smaller time-step size. If the oscillatory nature of f_c is due to the time discretization error in the evaluation of r and f_c in Equations (17) and (18), it should depend on the size of time stepping. However, the temporal variation of f_c in the smaller time-step size was almost the same as in Figure 1. We also checked the dependence on

the LTRAC parameter p . Note that as discussed in Appendix B, the larger p is, the larger f_c is expected to be. If the oscillatory nature of f_c is caused by the feedback between the broadening effect by the thermal conduction (enhanced by f_c) and the decrease in f_c (due to the smoothed temperature gradient), the period of the temporal variation of f_c would depend on the parameter p . However, the period of the oscillation did not show strong dependence on the parameter p . Therefore, the oscillatory nature of f_c may not be due to the thermal conduction or the temporal discretization, but to other processes associated with the thinning of the transition region, like the radiative cooling and/or acoustic waves (related to the hydrostatic equilibrium).

Figure 2 shows the spatial variation of the temperature, mass density, velocity, and the transition region broadening factor f_c at the snapshot of 800 s. Similar to the spatially averaged variables in Figure 1, the underresolved SH solution exhibits the largest difference from the reference solution, and the LTRAC solution is closest to the reference solution. The spatial profiles of the temperature and mass density obtained using the TRAC and LTRAC methods are close to the reference profile. In terms of the spatial profile of the velocity, both the TRAC and LTRAC solutions deviate from the reference solution by at least several kilometers. In addition, the velocity profiles from the TRAC and LTRAC solutions are not as smooth as that of the reference solution, which may be a numerical artifact. An important difference between the TRAC and LTRAC methods can be observed in the spatial profile of the broadening factor f_c . The value of f_c in the LTRAC solution increases only in the transition region, whereas that value in the TRAC solution increases all the way from the surface ($s = 0$) to the transition region. Considering the good agreement of the LTRAC solution to the reference solution, the increased broadening factor f_c in the upper chromosphere of the TRAC solution is not necessary to reproduce the dynamics and energetics of the coronal plasma. Moreover, the redundant enhancement of thermal conduction in the upper chromosphere may lead to weak chromospheric evaporation, which causes an excess of low-temperature emission, as shown in Section 3.2.

3.2. Effects on Spectroscopic Observables

To evaluate the performance of the LTRAC method on spectroscopic observables, we used the velocity DEM described in Section 2.3. Figure 3 shows the quantities calculated from the VDEM and temporally averaged from 100 to 1500 s. The DEM from the underresolved SH solution shows the largest deviation from the reference solution from the transition region temperature to the coronal temperature. From the resolution dependence within the range of $6.25 \text{ km} \leq \Delta s \leq 100 \text{ km}$, we found that the underresolved SH solution always overestimates (underestimates) the DEM in $T < 10^{5.2} \text{ K}$ ($T > 10^{5.2} \text{ K}$). The DEM from the TRAC solution is close to that from the reference solution in $T > 10^{5.4} \text{ K}$. However, the TRAC solution overestimates the DEM in lower temperatures of $T < 10^{5.4} \text{ K}$. This excess of low-temperature emission in the TRAC solution was observed even when a finer grid size was used. We speculate that enhanced thermal conduction in the upper chromosphere to the lower transition region in the TRAC solution may be heating the high-density, cool plasma in the lower layers and enhancing the emission measure in the upper transition region. In contrast, the LTRAC method successfully reproduces the DEM in the wider

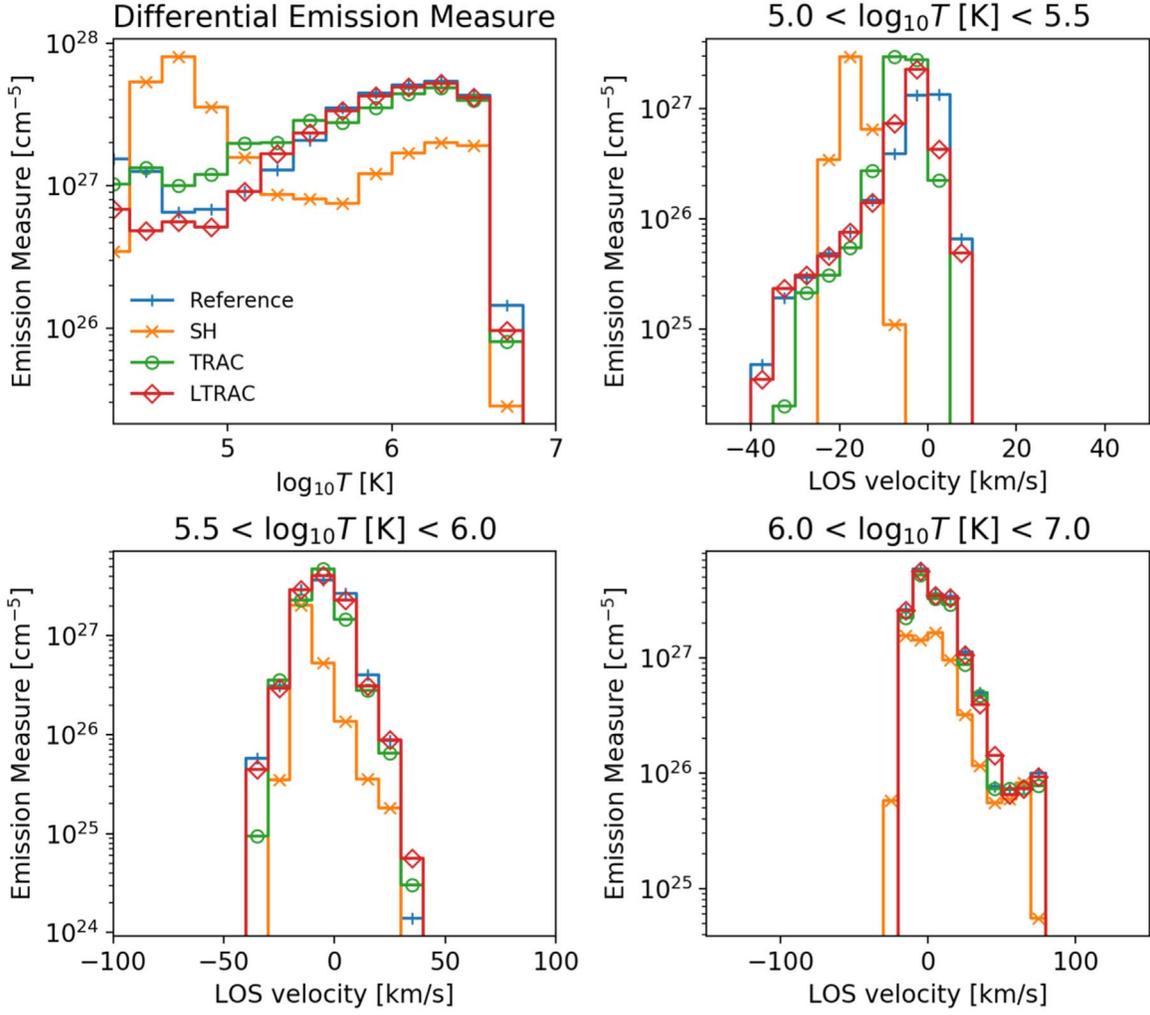


Figure 3. Time-averaged DEM (top left) and the velocity DEM averaged over the temperature intervals of $[10^{5.0}, 10^{5.5}]$ K (top right), $[10^{5.5}, 10^{6.0}]$ K (bottom left), and $[10^{6.0}, 10^{7.0}]$ K (bottom right). The notation is the same as that in Figure 1.

temperature range of $T > 10^{5.0}$ K. This result indicates the advantage of the localized broadening factor used in the LTRAC method. In the lowermost temperature of $T < 10^{5.0}$ K, a slight lack of the emission measure can be observed. The LTRAC solution was found to show smaller DEM even in the higher resolution (down to $\Delta s = 6.25$ km) in this temperature range. We also note that the lack/excess of the emission measure in $T < 10^{5.0}$ K depends on the parameter p in the LTRAC method (see Appendix B).

We averaged the velocity DEM over the temperature intervals $[10^{5.0}, 10^{5.5}]$ K, $[10^{5.5}, 10^{6.0}]$ K, and $[10^{6.0}, 10^{7.0}]$ K to mimic a line profile produced by the optically thin emission. The resulting profile differs from the actual line profile (calculated using the coronal approximation) because the temperature-averaged VDEM profile does not include broadening processes (e.g., thermal or instrumental broadenings). However, we believe that at least a part of the effect of the LTRAC method on spectroscopic observables can be investigated using the VDEM. The velocity dependence of the VDEM from the underresolved SH solution deviates from the reference solution, especially in the lower-temperature intervals. The VDEM profile from the TRAC solution agrees well for $T > 10^{5.5}$ K with the reference profile, but it deviates slightly at lower temperature. In contrast, the VDEM profile calculated

from the LTRAC solution shows better agreement with that of the reference solution, especially in the low-temperature interval.

To evaluate the differences in the VDEM more quantitatively, we calculated the time variation of the velocity moments of the VDEM as shown in Figure 4. The velocity shift V_{shift} and velocity width V_{width} for the temperature interval $T_0 < T < T_1$ are defined as

$$V_{\text{shift}} = \frac{\int_{T_0}^{T_1} \int_V V \cdot \text{VDEM}(T, V) dV dT}{\int_{T_0}^{T_1} \int_V \text{VDEM}(T, V) dV dT} \quad (24)$$

and

$$V_{\text{width}} = \sqrt{\frac{\int_{T_0}^{T_1} \int_V V^2 \text{VDEM}(T, V) dV dT}{\int_{T_0}^{T_1} \int_V \text{VDEM}(T, V) dV dT}}, \quad (25)$$

respectively. All of the SH, TRAC, and LTRAC solutions produce oscillatory behavior in the lower-temperature interval, because fewer grid points are used for calculating the VDEM in the transition region. The temporal oscillation of the underresolved SH solution is the most significant because the transition region is

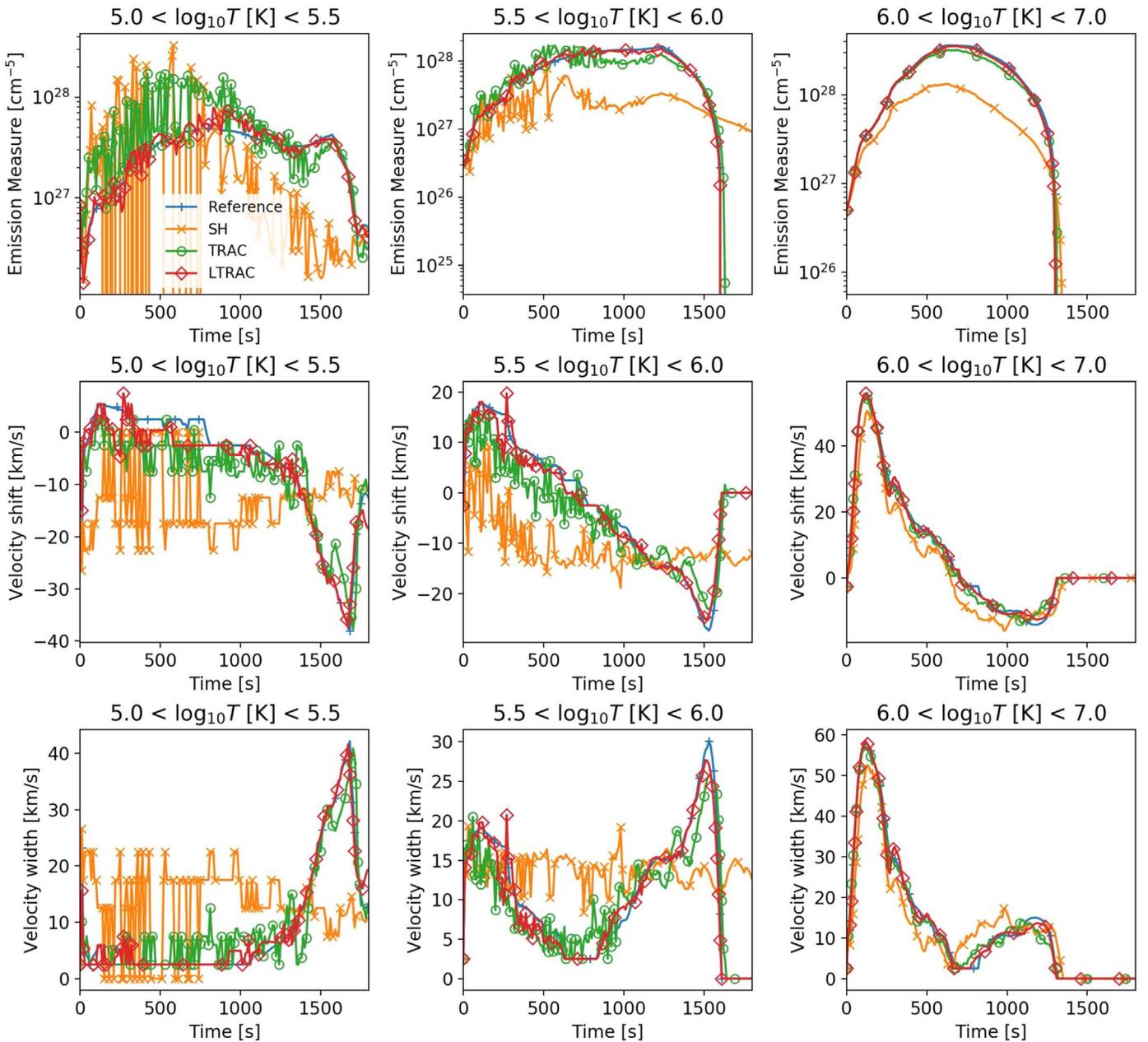


Figure 4. Time variation of the emission measure (top row), the velocity shift (middle row), and the velocity width (bottom row), calculated from the VDEM averaged over the temperature intervals of $[10^{5.0}, 10^{5.5}]$ K (left column), $[10^{5.5}, 10^{6.0}]$ K (middle column), and $[10^{6.0}, 10^{7.0}]$ K (right column). The notation is the same as that in Figure 1.

resolved by only a few grid points in the numerical solution. As the TRAC and LTRAC solutions both have more grid points in the transition region, the temporal oscillation is slightly suppressed using these two methods. Among the three methods, we found that the LTRAC solution exhibits the smallest temporal oscillation and provides the best agreement with the reference solution. The LTRAC solution can reproduce the velocity shift and width within the difference of several kilometers per second to the reference solution even with a coarse grid size of 50 km.

3.3. Convergence Analysis

Figure 5 shows a convergence analysis of the spatially averaged variables (shown in Figure 1). The error is defined as the mean absolute difference from the reference solution. The

convergence rate of the SH method is very small (less than first order with respect to the grid size Δs). The TRAC method shows similarly slow convergence, although the absolute value of the error is much smaller than that of the SH method. The LTRAC method produces slightly faster convergence of the temperature at the loop top and of the velocity at the footpoint. The mass density at the loop top in the LTRAC method displays irregular convergence for $\Delta s \leq 25$ km. We speculate that the mass density of the LTRAC simulation with $\Delta s = 25$ km accidentally approached the value in the reference solution. In the high-resolution range of $\Delta s \leq 12.5$ km, both the mass density and the velocity of the LTRAC solutions exhibit slow convergence, similar to the SH and TRAC methods. These results indicate that although the TRAC and LTRAC methods both yield smaller deviations from the

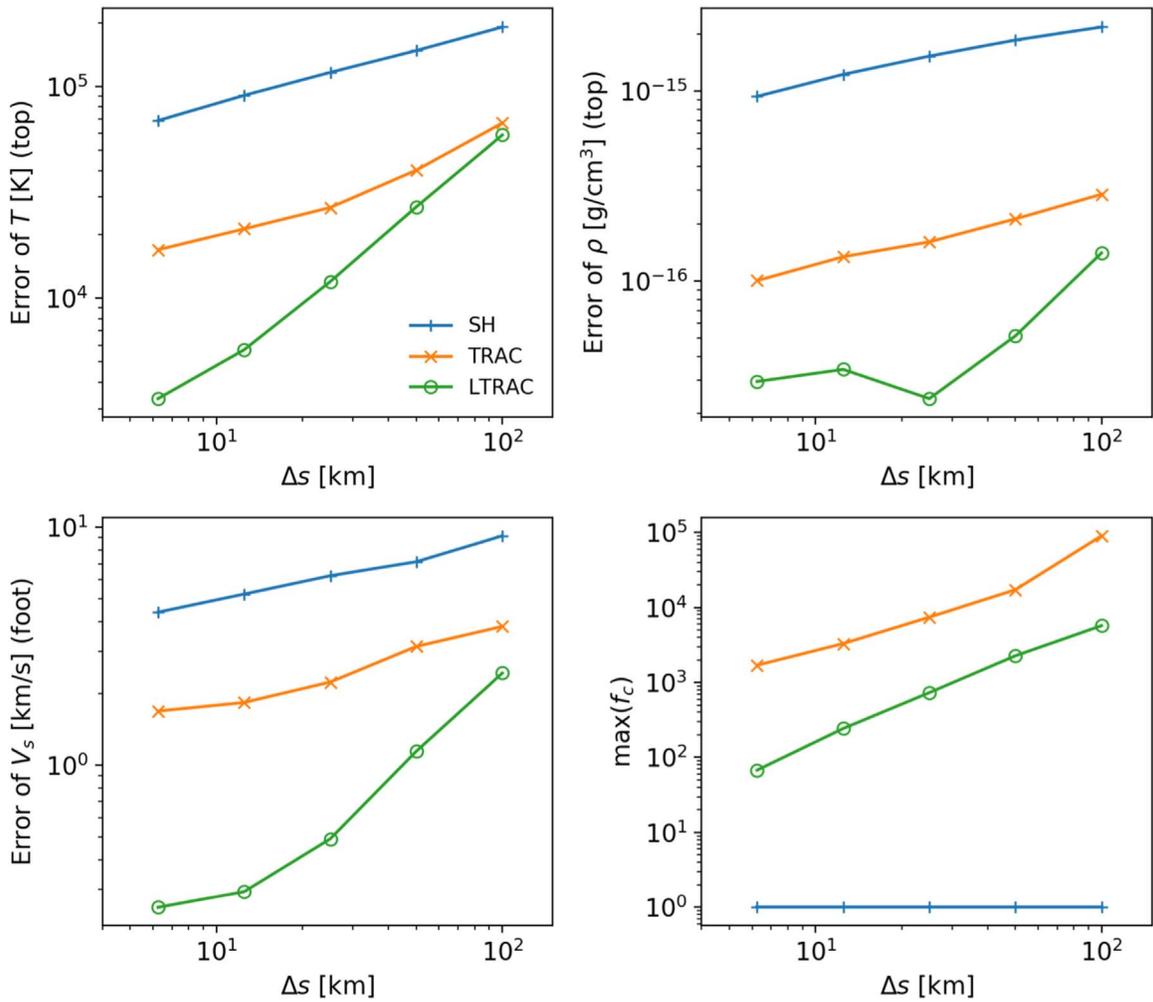


Figure 5. Convergence of the spatially averaged variables as a function of the grid size. Shown are the mean absolute errors measured by the temperature at the loop top (top left), the mass density at the loop top (top right), and the velocity at the footpoint (bottom left), as well as the maximum value of the broadening factor f_c (bottom right). The results for the SH method (blue line with plus signs), the TRAC method (orange line with crosses), and the LTRAC method (green line with circles) are shown.

reference solution, the rate of convergence in the high-resolution simulation may not be significantly improved from the SH method. We also measured the dependence of the broadening factor f_c on the grid size by the time average of the spatial maximum. The values of the broadening factor f_c in the TRAC and LTRAC solutions show approximately a first-order convergence rate.

We also analyzed the convergence of the emission measure, the velocity shift, and the velocity width as shown in Figure 6. The LTRAC method shows the smallest error in the emission measure among the three methods, in all three temperature intervals: $[10^{5.0}, 10^{5.5}]$ K, $[10^{5.5}, 10^{6.0}]$ K, and $[10^{6.0}, 10^{7.0}]$ K. The convergence of the emission measure is slow, which may be due to the slow convergence of the mass density shown in Figure 5. The velocity shift and width show better convergence rates (especially in the LTRAC solutions), but the grid convergence is less than first order. The errors in the velocity shift and width in the LTRAC solutions are less than 2 km s^{-1} even when a coarse grid size of 50 km is used. These results suggest that the Doppler shift and nonthermal line width of the synthesized optically thin line emission from the LTRAC solution may be accurate enough to be compared with spectroscopic observations.

4. Discussion

From the nature of the transition region, a large number of grid points are required to resolve this thin layer fully. In this study, we have proposed a new numerical method, called the LTRAC method, which enables the physically accurate coronal dynamics and energetics to be obtained with a coarse grid size. Following the strategy of Lionello et al. (2009) and Johnston & Bradshaw (2019), the LTRAC method broadens the unresolved transition region by modifying the thermal conduction coefficient and the radiative cooling rate. The major difference between the LTRAC method and the previous methods is that the broadening factor of the transition region is concentrated only in the region with the high-temperature gradient. The localized broadening factor provides better reproduction of the emission measure especially in the lower-temperature region, which may result from the suppression of excess conductive heating in the upper chromosphere observed in the previous methods. We investigated the possible effect of the Doppler velocity and the nonthermal broadening of optically thin emission lines based on the velocity DEM from the numerical solutions. We found that the synthesized profile of the optically thin line emission is reproduced within an error of several kilometers per second with a coarse grid size of 50 km. The LTRAC method is designed so that it can be extended

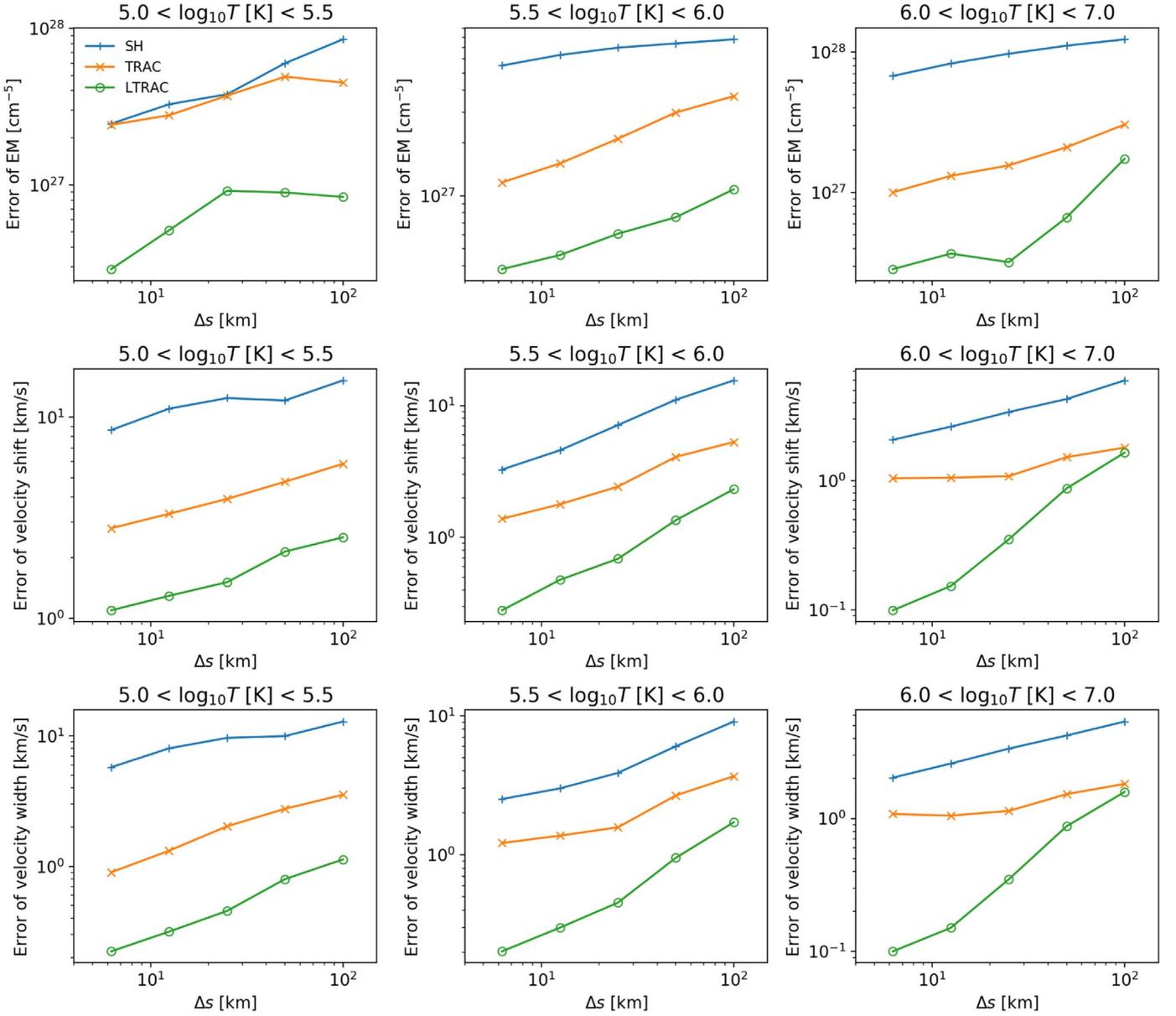


Figure 6. Convergence of the VDEM against the grid size. Shown are the mean relative error of the emission measure (top row), the mean error of the velocity shift (middle row), and the mean absolute error of the velocity width (bottom row), calculated from the VDEM averaged over the temperature intervals of $[10^{5.0}, 10^{5.5}]$ K (left column), $[10^{5.5}, 10^{6.0}]$ K (middle column), and $[10^{6.0}, 10^{7.0}]$ K (right column). The notation is the same as that in Figure 5.

to multidimensional simulations using the domain decomposition technique. Application of the LTRAC method to multidimensional coronal heating models (e.g., Abett 2007; Gudiksen et al. 2011; Iijima & Yokoyama 2015, 2017; Rempel 2017) will allow us to reproduce the wave reflection rate at the transition region and the energy loss from the corona more accurately with a coarse grid spacing.

In this paper, we evaluated the performance of the proposed method using one-dimensional hydrodynamic simulations of a coronal loop in an active region. In flaring loops with higher temperatures, the accuracy of the coarse-resolution simulations may be worse than that for cooler loops. It should be noted that we assumed a uniform volumetric heating rate in this study. The performance of the proposed method should be tested with a wider range of numerical settings, including a spatially nonuniform heating rate. However, we believe that the proposed method will be effective even in such problems, as

the free parameters in the LTRAC method determine the sensitivity based only on the normalized temperature gradient and are not chosen for specific problems.

We have ignored the realistic physical processes such as the nonequilibrium ionization or the effect of optical thickness in this study. The numerically broadened transition region may require special care on these processes. For example, in calculating the VDEM (Section 2.3), we accounted for the broadening factor f_c by reducing the effective line-of-sight distance by a factor of $1/f_c$ (see Equation (23)), which hopefully conserves the total emission measure from the artificially broadened region. Similarly, in calculating the optical thickness, the line-of-sight distance must be modified to account for the broadening factor. Modification nonequilibrium ionization processes may be also required. If we assume that an ion moves across the broadened transition region at a constant velocity, the total number of radiative/collisional reactions the ion

undergoes may increase by a factor of f_c . To prevent this artifact, the reaction rates should also be reduced by the broadening factor f_c . We emphasize that these suggestions are mere speculations, and they should be verified more precisely in a future study.

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Appendix A Convergence of the Reference Solution

As the requirement on the grid size is extremely severe in the transition region, we were not able to resolve the transition region fully. Figure 7 compares high-resolution runs obtained by the SH, TRAC, and LTRAC methods. The nonuniform grid

spacing is used with a grid size of 100 m near the transition region. The variations of the temperature, mass density, and velocity field calculated by these three methods are very close to each other, which implies that the numerical solutions are roughly converged at this grid size. If the transition region in the numerical solution is fully resolved, the normalized temperature gradient of Equation (17) should be smaller than the critical value ($\exp(\sigma_c)$ for the TRAC method and r_c for the LTRAC method), which implies that the maximum value of the f_c should be unity in both the TRAC and LTRAC methods. The nearly constant (but larger than unity) value of f_c in the TRAC solution after 1000 s occurs because the lower limit of T_c is set to 2×10^4 K. Unfortunately, $\max(f_c)$ in the LTRAC solution is about 5, which indicates that the numerical grid size should be 20 m or less to obtain a truly converged solution (in the sense that the LTRAC solution becomes exactly identical to the SH solution). Such a high-resolution simulation is difficult to achieve even in a one-dimensional domain, as we do not use the adaptive mesh refinement. This requirement that the grid size be less than 20 m is consistent with the estimate obtained using the static loop model of Rosner et al. (1978), which predicts that the thickness of the transition region is about

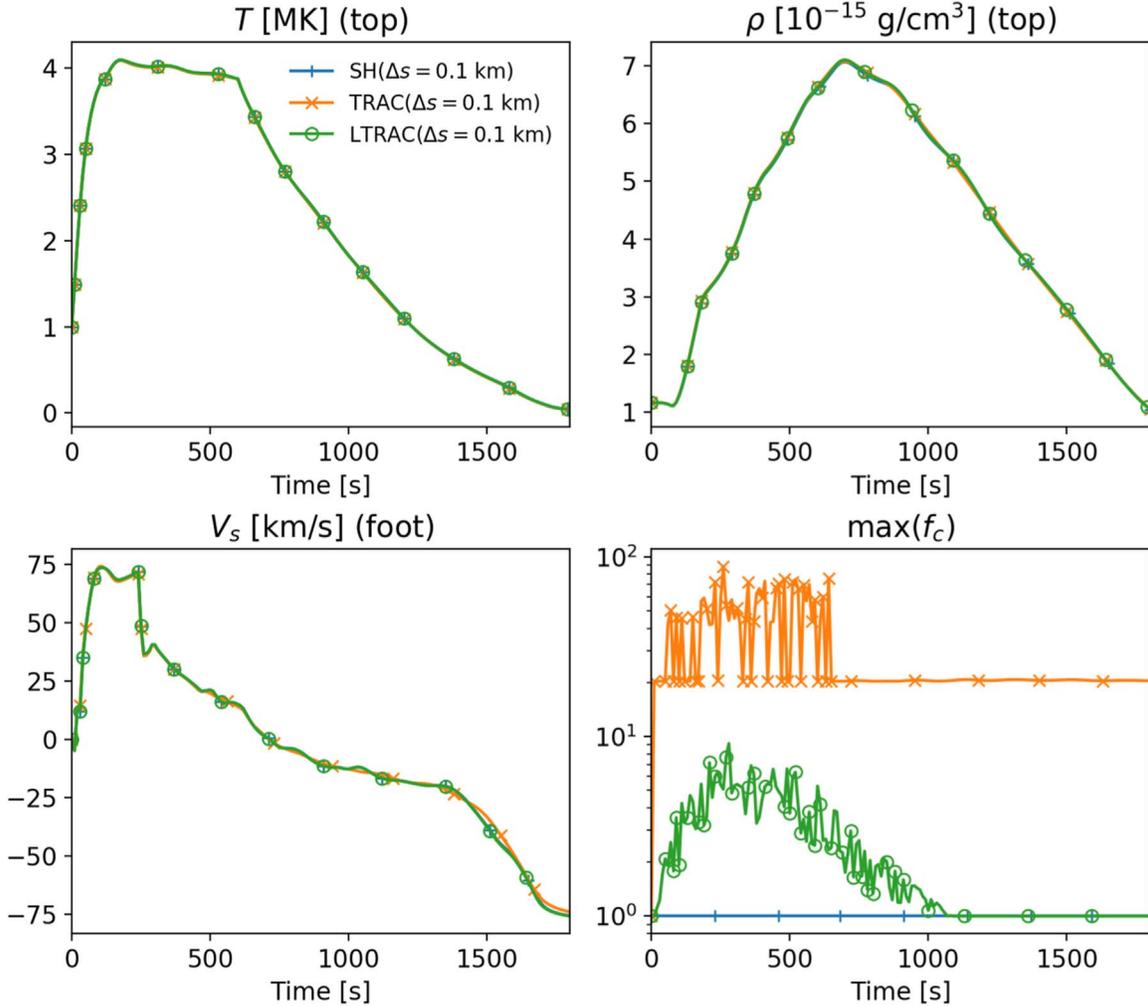


Figure 7. Time variation of the temperature at the loop top (top left), the mass density at the loop top (top right), the velocity at the footpoint (bottom left), and the maximum value of the broadening factor f_c (bottom right). Shown are the numerical solutions calculated using the SH method (orange line with crosses), using the TRAC method (green line with circles), and using the LTRAC method (red line with diamonds). The nonuniform grid spacing is used with the grid size of 100 m near the transition region.

400 m for a loop top temperature of 4 MK (see Section 1). The results presented in this paper do not depend on the choice of the reference solution, except for the convergence rates of the highest-resolution runs (grid sizes less than 20 km) in Appendix B. In this study, we use the SH solution with a grid size of 100 m as the reference solution.

Appendix B Dependence on LTRAC Parameters

The LTRAC method employs two parameters, the critical value r_c of the temperature gradient and the acceleration factor p . Here, we briefly investigate the dependence of the LTRAC solutions on these parameters. Figure 8 shows the dependence of the convergence rates on the LTRAC parameter p , fixing the value of r_c to be 1.5. Clearly, a larger value of p produces faster grid convergence. The error in the LTRAC method with $p = 2$ is similar to that of the TRAC method (Figure 5).

Figure 9 shows the dependence of the convergence rate on r_c , with the value of p fixed to 4.0. The result is not very sensitive to the parameter r_c , although a smaller r_c tends to produce smaller errors and larger values of the broadening factor f_c . The reader may notice that the mass density of

$r_c = 1.2$ apparently converges to the reference value regularly. However, we note that this regular convergence may be produced by the insufficient resolution of the reference solution. If we use the high-resolution LTRAC run as a reference, the mass density shows irregular convergence in all values of r_c . We chose the typical value $r_c = 1.5$ so that the threshold of the temperature gradient would be close to the value in the TRAC method (so that $\exp(\sigma_c) \sim r_c$).

In Section 2.3, we have mentioned a small lack of the low-temperature DEM in $T < 10^{5.0}$ K found in the LTRAC solution. The lack and excess of the DEM in this low temperature range depend on the parameter p . Figure 10 shows the parameter dependence of the DEM. The uniform grid size of 50 km is used except the reference solution. In the cases of smaller p -value (i.e., $p \leq 2$), the emission measure in $T < 10^{5.0}$ K tends to show slight excess, in contrast to the cases for $p \geq 3$. This dependence on the parameter p was observed independently of the spatial resolution (down to $\Delta s = 6.25$ km). This result may imply that the optimal parameter of p lies between 2 and 3 in terms of the DEM reproducibility. We found that the dependence of the low-temperature DEM on the parameter r_c is not significant.

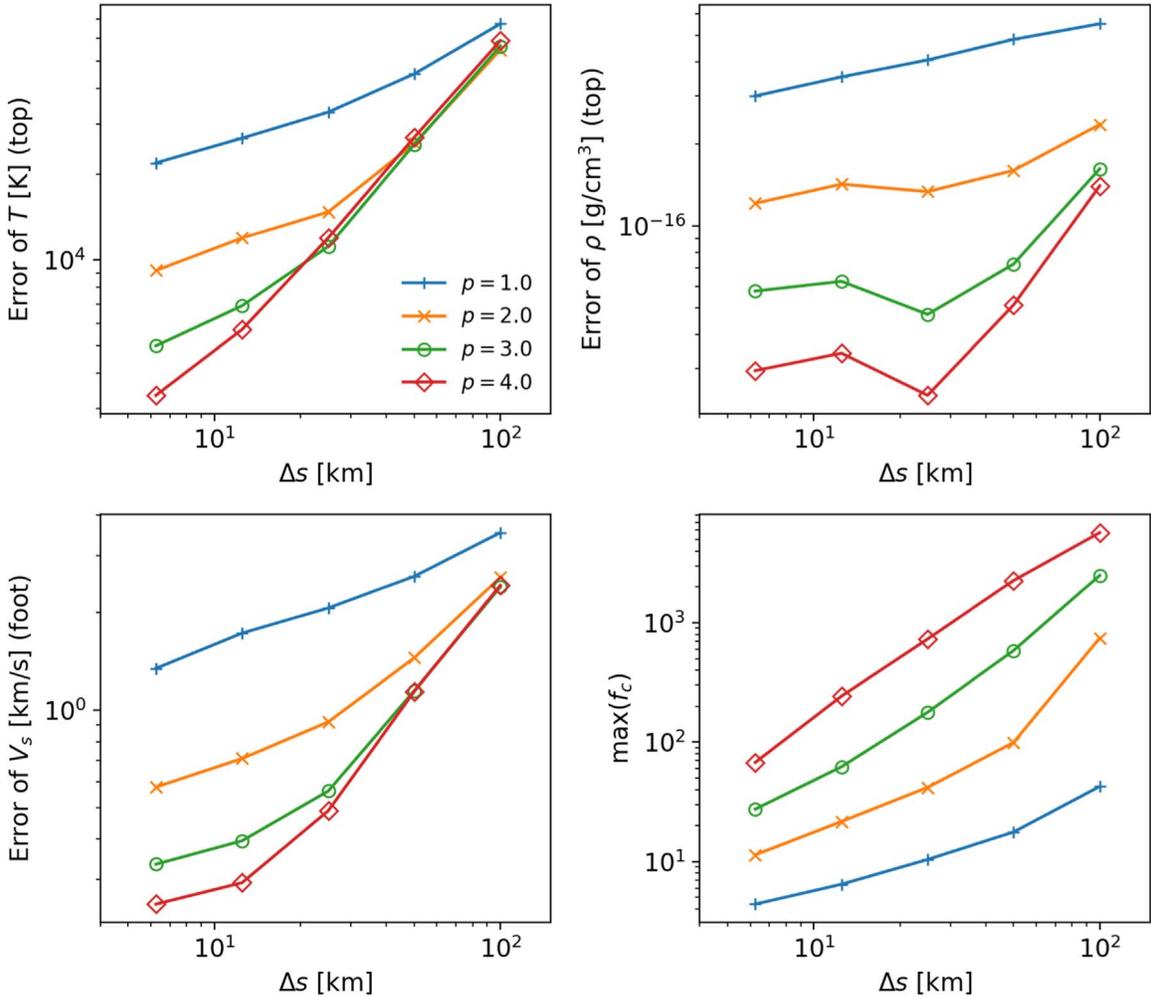


Figure 8. Same as Figure 5, but showing the dependence on the LTRAC parameter p . The results for $p = 1$ (blue line with plus signs), $p = 2$ (orange line with crosses), $p = 3$ (green line with circles), and $p = 4$ (red line with diamonds) are shown.

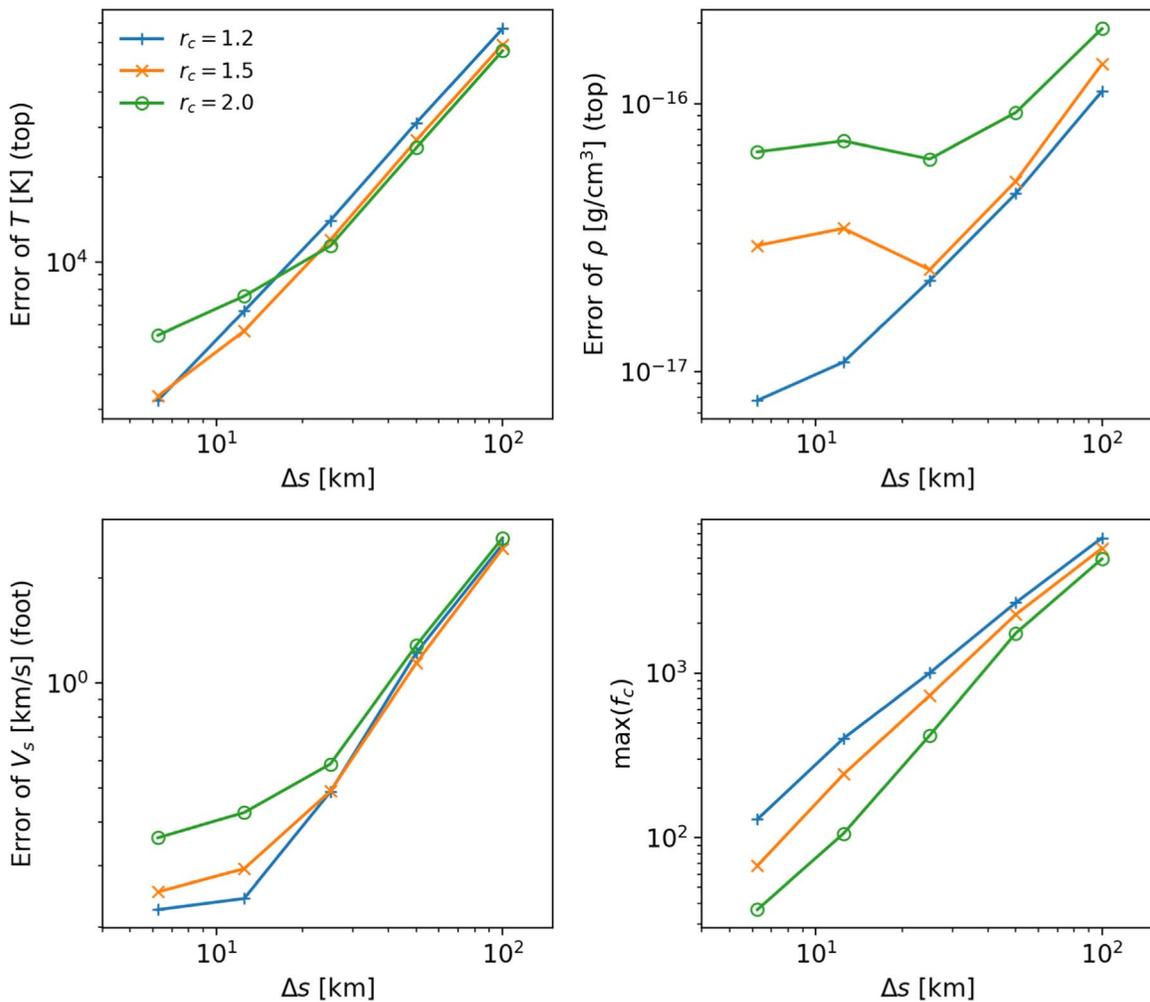


Figure 9. Same as Figure 5, but showing the dependence on the LTRAC parameter r_c . The results for $r_c = 1.2$ (blue line with plus signs), $r_c = 1.5$ (orange line with crosses), and $r_c = 2$ (green line with circles) are shown.

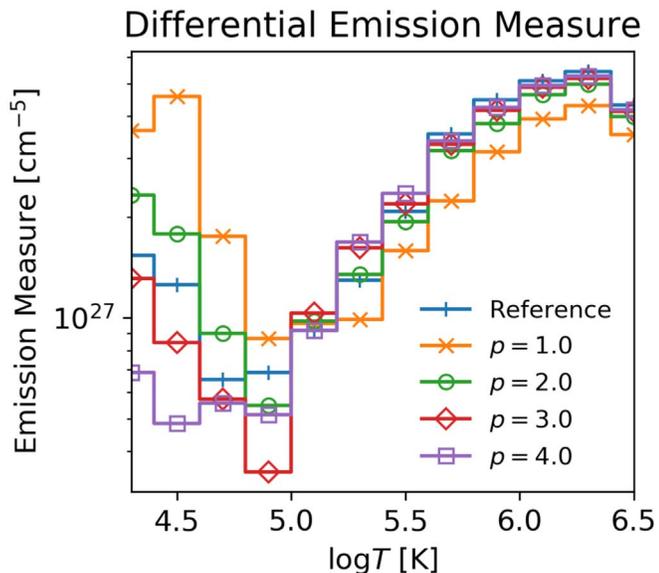


Figure 10. Dependence of the time-averaged DEM on the LTRAC parameter p . The results for $p = 1$ (orange line with crosses), $p = 2$ (green line with circles), $p = 3$ (red line with diamonds), $p = 4$ (purple line with boxes), and the reference (blue line with plus signs) are shown.

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